2010

Distance to U Pegasi by the DDE Algorithm

Todd R. Vaccaro
St. Cloud State University, trvacarro@stcloudstate.edu

Dirk Terrell
Southwest Research Institute

Robert E. Wilson
University of Florida

Follow this and additional works at: https://repository.stcloudstate.edu/phys_facpubs
Part of the Stars, Interstellar Medium and the Galaxy Commons

Recommended Citation
Vaccaro, Todd R.; Terrell, Dirk; and Wilson, Robert E., "Distance to U Pegasi by the DDE Algorithm" (2010). Physics and Astronomy Faculty Publications. 2.
https://repository.stcloudstate.edu/phys_facpubs/2

This Article is brought to you for free and open access by the Department of Physics and Astronomy at theRepository at St. Cloud State. It has been accepted for inclusion in Physics and Astronomy Faculty Publications by an authorized administrator of theRepository at St. Cloud State. For more information, please contact modea@stcloudstate.edu, rsweelbaum@stcloudstate.edu.
Distance to U Pegasi by the DDE Algorithm

T. R. Vaccaro

Francis Marion University, Dept. of Physics & Astronomy, P.O. Box 100547, Florence, SC 29502, USA

D. Terrell

Southwest Research Institute, Dept. of Space Studies, Boulder, CO 80301, USA

R. E. Wilson

University of Florida, Dept. of Astronomy, Bryant Space Science Center, P.O. Box 112055, Gainesville, FL 32611-2055, USA

Abstract. A distance is found for the W UMa type binary U Pegasi, with a newly modified version of the Wilson-Devinney program (W-D) that makes use of the direct distance estimation (DDE) algorithm. The reported distance of $d = 123.6\text{pc}$ is an average based on solutions for $B$ and $V$ data and a primary star temperature of 5800K. Standardized light curves (not differential), radial velocities, and a spectroscopic primary star temperature are input to the program. Differential corrections were performed for each light curve band along with the velocities for two primary temperatures that span 100K. $\log_{10}d$ is a model parameter like many others that are adjustable in W-D. The eclipsing binary distance agrees with the Hipparcos parallax distance and is more precise.

1. Analysis

Light curves of the W-type W UMa eclipsing binary U Pegasi by Zhai et al. (1984) were converted from differential magnitudes to the standard $B$ and $V$ systems by addition of the comparison star’s magnitudes, $B = 10^m.092 \pm 0^m.026$, $V = 9^m.488 \pm 0^m.021$ (Kharchenko 2001) for BD+14°5078. Both U Peg components are of type G2 (Lu 1985), and an adopted primary star temperature of 5800K was rounded from two values in Cox (2000). Subsequent work by Zhai et al. (1988) refined several velocity points that occur near conjunction. The light/velocity data were solved simultaneously for distance and other parameters by a version of the Wilson-Devinney (W-D) Differential Corrections (DC) program that contains the direct distance estimation (DDE) algorithm.

The Zhai et al. (1988) parameters were initial input. However the linear ephemeris (Zhai et al. 1984) was independently adjusted to improve fits to $V$ band light curves and velocity curves, which are separated by 6 years. A solution with primary temperature 5900K was also made to allow parameter interpolation in $T_1$ by readers. Adjusted parameters were inclination ($i$), secondary temperature ($T_2$), semi-major axis ($a$), potential ($\Omega_1$), systemic velocity ($V_y$), mass ratio ($q = M_2/M_1$), and distance ($\log_{10}d$). Initial solutions that adjusted third light found $\ell_3$ to be insignificant, so $\ell_3$ was fixed at zero. The parameters were split into two DC subsets to reach the solutions of Table 1.
2. Final Remarks

The theory of the DDE algorithm, with overviews of simulations, is in Wilson (2008). As applied to 1978 Zhai, et al. light curves and 1984 Lu velocity data for U Peg, DDE yields respective $B, V$ distances of 121.6 and 125.5 pc if $T_1 = 5800K$. The formal errors are smaller than the spread, so calibration errors probably account for most of the modest disagreement. Hipparcos parallaxes, as revised by van Leeuwen (2007), span 122 to 175 pc according to the reported 1σ errors. Our model light curves might fit the data better if spots were modeled as in Djurašvić et al. (2001), but derived distance typically is not much affected by such refinement. Although period variations have been linked to a third body by several authors, our solutions indicate that $\ell_3$ is unlikely to affect distance estimates significantly. Another set of distances will be based on light curves by T. Pribulla (Pribulla & Vaňko 2002) and published elsewhere.

This work was supported in part by NSF grant AST-0307561 at the University of Florida.

Table 1. Solutions for U Peg

<table>
<thead>
<tr>
<th>Parameter</th>
<th>$B(T_1 = 5800K)$</th>
<th>$V(T_1 = 5800K)$</th>
<th>$B(T_1 = 5900K)$</th>
<th>$V(T_1 = 5900K)$</th>
</tr>
</thead>
<tbody>
<tr>
<td>$d[\text{''}]$</td>
<td>75.26 ± 0.08</td>
<td>75.44 ± 0.08</td>
<td>74.92 ± 0.09</td>
<td>75.35 ± 0.12</td>
</tr>
<tr>
<td>$T_2[\text{K}]$</td>
<td>5637 ± 4</td>
<td>5622 ± 8</td>
<td>5738 ± 8</td>
<td>5619 ± 7</td>
</tr>
<tr>
<td>$a[R_\odot]$</td>
<td>2.56 ± 0.02</td>
<td>2.54 ± 0.02</td>
<td>2.57 ± 0.02</td>
<td>2.56 ± 0.02</td>
</tr>
<tr>
<td>$v_{\gamma}$ [km s$^{-1}$]</td>
<td>30.9 ± 0.8</td>
<td>30.7 ± 1.0</td>
<td>31.0 ± 0.9</td>
<td>31.2 ± 0.9</td>
</tr>
<tr>
<td>$\Omega_1$</td>
<td>6.294 ± 0.012</td>
<td>6.207 ± 0.025</td>
<td>6.256 ± 0.009</td>
<td>6.201 ± 0.013</td>
</tr>
<tr>
<td>$q = M_2/M_1$</td>
<td>2.839 ± 0.009</td>
<td>2.763 ± 0.019</td>
<td>2.824 ± 0.006</td>
<td>2.772 ± 0.010</td>
</tr>
<tr>
<td>$d[\text{pc}]$</td>
<td>121.6 ± 1.0</td>
<td>125.5 ± 1.0</td>
<td>129.3 ± 1.0</td>
<td>132.2 ± 1.0</td>
</tr>
<tr>
<td>$R_1/a$ (pole)</td>
<td>0.2809 ± 0.0011</td>
<td>0.2819 ± 0.0023</td>
<td>0.2826 ± 0.0008</td>
<td>0.2830 ± 0.0012</td>
</tr>
<tr>
<td>$R_1/a$ (side)</td>
<td>0.2938 ± 0.0013</td>
<td>0.2947 ± 0.0028</td>
<td>0.2958 ± 0.0010</td>
<td>0.2961 ± 0.0015</td>
</tr>
<tr>
<td>$R_1/a$ (back)</td>
<td>0.3327 ± 0.0050</td>
<td>0.3327 ± 0.0050</td>
<td>0.3358 ± 0.0018</td>
<td>0.3352 ± 0.0027</td>
</tr>
<tr>
<td>$R_2/a$ (pole)</td>
<td>0.4506 ± 0.0008</td>
<td>0.4474 ± 0.0018</td>
<td>0.4514 ± 0.0006</td>
<td>0.4490 ± 0.0010</td>
</tr>
<tr>
<td>$R_2/a$ (side)</td>
<td>0.4483 ± 0.0012</td>
<td>0.4803 ± 0.0025</td>
<td>0.4854 ± 0.0009</td>
<td>0.4824 ± 0.0013</td>
</tr>
<tr>
<td>$R_2/a$ (back)</td>
<td>0.5131 ± 0.0015</td>
<td>0.5089 ± 0.0032</td>
<td>0.5148 ± 0.0011</td>
<td>0.5114 ± 0.0017</td>
</tr>
</tbody>
</table>

References

Kharchenko, N. V. 2001, Kinematika Fiz. Nebesnykh Tel, 17, 409
Lu, W. X. 1985, PASP, 97, 1086
Pribulla, T., & Vaňko, M. 2002, CoSka, 32, 79